

Field Demonstration of Coherent Uplink from a Phased Array of Widely Separated Antennas: Steps Toward A Verifiable Real-Time Atmospheric Phase Fluctuation Correction for a High Resolution Radar System

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ABSTRACT

With the successful field demonstration of uplink arraying at 8 GHz (X-band) at NASA's Goldstone tracking station in California and at 14 GHz (Ku-band) near the Jet Propulsion Laboratory in California, NASA is pursuing a similar demonstration of the capability at 7.1 GHz (X-band) and 30-31 GHz (Ka band) and incorporating real-time correction for atmospheric phase fluctuations. Such a demonstration would then enable NASA to establish a high power, high resolution, 24/7 availability radar system (a) for tracking and characterizing observations of Near Earth Objects, (b) tracking, characterizing and determining the statistics of small-scale (≤ 10 cm) orbital debris, (c) to incorporate the capability into its space communication and navigation tracking stations for emergency spacecraft commanding in the Ka band era which NASA is entering, and (d) to field capabilities of interest to other US government agencies. We are beginning the demonstration in a communications mode to more easily verify the phased array combining and real-time atmospheric compensation aspects of the system. Additionally, we are taking a two-pronged demonstration approach: (a) at X-band at NASA's Goldstone tracking facility using 34m antennas of the Deep Space Network (DSN) targeting NASA spacecraft, and (b) at the Kennedy Space Center first at X-band and then at Ka band targeting both commercial and military satellites. In this paper, we present the steps we are undertaking to produce the first verifiable phased array combining at X-band using widely separated (~60 m) antennas- including *the first ever successful uplink arraying to a satellite orbiting another planet*. We then describe our visionary program to achieve similar results at Ka band in both communication and, notionally, radar applications: Evolutionary Steps Leading to Revolutionary Increases in Capability and Capacity.

1.0 INTRODUCTION

NASA has embarked on a path to implement a high power, higher resolution radar system to better track and characterize near Earth objects [NEO's] and orbital debris. We are advancing an X/Ka band system (KaBOOM: Ka Band Objects Observation and Monitoring) to supplement the S-band radar at Arecibo, Puerto Rico and the X-band radar at NASA's Goldstone tracking complex in California. The three facilities will complement each other in that different wavelengths have different resolutions and penetration depths. The X/Ka band radar system also has applications for cost effective space situational awareness. This work describes our path toward demonstrating Ka band coherent uplink arraying with real-time atmospheric compensation using three 12m diameter antennas at the Kennedy Space Center. Coherent uplink arraying has been successfully demonstrated by three NASA groups: at X band and at Ku band, without and perhaps atmospheric compensation, and by sending commands to and receiving telemetry from GEO and deep space satellites. Deep space in NASA terms means a distance greater than 2 million kilometers.

KaBOOM is a Ka band coherent uplink arraying proof of concept demonstration being undertaken to allow decisions to be made for implementing a National Radar Facility [large scale array(s)]:

- High power, high resolution radar system
- Space Situational Awareness
- 24/7 availability for NEO and orbital debris tracking and characterization
- Map out radar stealth zones on Mars- help define "no drive" zones for future rovers to avoid the Spirit problem
- Beam sailing propulsion capability

2.0 THE ADVANTAGES OF SITE DIVERSE, MULTI-FREQUENCY RADAR OBSERVATIONS OF NEAR EARTH OBJECTS

NASA has several major uses for uplink arraying: (1) planetary defense- tracking and characterization (size, shape, spin, porosity, surface features, non-gravitational forces, presence of moons, etc.) of near Earth objects (NEOs), (2) improved detection/tracking of small (≤ 10 cm) orbital debris particles, (3) rapidly available high power emergency uplink capability for spacecraft emergencies, and (4) radio science experiments (tomography of planetary atmospheres, general relativity tests, mass determinations, occultations, surface scattering, etc.).

The NASA Authorization Act of 2008 directs NASA to catalog 90% of NEOs > 140 m in size. The Goldstone radar on the 70m antenna is typically available only 2-3% of the time due to spacecraft tracking obligations. It is further limited by the high power density of the beam since the transmitter is 450 kW which requires consultation and compliance with some 31 federal, state, local, and military authorities for permission to trigger the radar. It can take up to a month to obtain the necessary permission. In times of spacecraft emergencies, the time can be cut back to a few days, but it would be better if the radar system was available on demand as needed.

Tracking of orbital debris particles less than 10cm in size is difficult. Where high resolution is possible, the antenna beam is too small for accurate long term tracking; short term detection of individual particles is possible and the statistics of small particles are obtained, however. In those facilities where particle tracking is possible, the resolution is lacking.

Having a site and frequency diverse radar system is an advantage over today's Goldstone Solar System Radar. If NASA has both Goldstone and the KaBOOM site radar available, we can obtain longer tracks leading to improved NEO orbit determination. In addition, longer tracks will also allow for more precise data to determine the tumbling parameters which in turn lead to better models of the moments of inertia of the NEO. Since the KaBOOM site operates at X and Ka bands, we shall obtain higher range resolution than at Goldstone [5-10 cm vs. 400 cm] and, hence, produce sharper images of the NEO surface structures. By adding in Arecibo radar data at S-band (2 GHz), we could have a three frequency radar system, and since each frequency has a different penetration depth, NASA will obtain a greater knowledge of the NEO porosity parameters. All of these aspects are important for designing a multistatic radar system- essentially a very long baseline interferometry (VLBI) system- and carrying out crewed missions to NEOs whether they are asteroids or comets, etc.

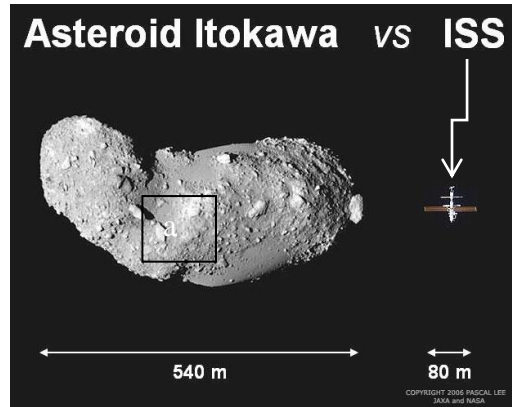


Fig 1. Relative sizes of asteroid Itokawa and the International Space Station. Note the boulders on the asteroid. For robotic and crewed missions to asteroids, NASA will need to know the surface structure - such as the boulders in the box - so as to ensure the safety of the landing spacecraft and its inhabitants.

An interesting capability that uplink arraying offers an opportunity to get high EIRP (effective isotropic radiated power) from relatively lower power transmitters because the uplink power, for identical transmitters and antennas, is proportional to N^2 , where N is the number of antennas in the array. The power is proportional to the number of transmitters (N) times the number of antennas (N); hence, N^2 . Figure 2 demonstrates this effect for data taken of the planet Venus with two of the 34m antennas at Goldstone. Figure 2a shows the radar image for a single 34m antenna. Figure 2b shows the radar image, with a 6 dB gain (factor of 4) when the power from two 34m dishes are coherently combined.

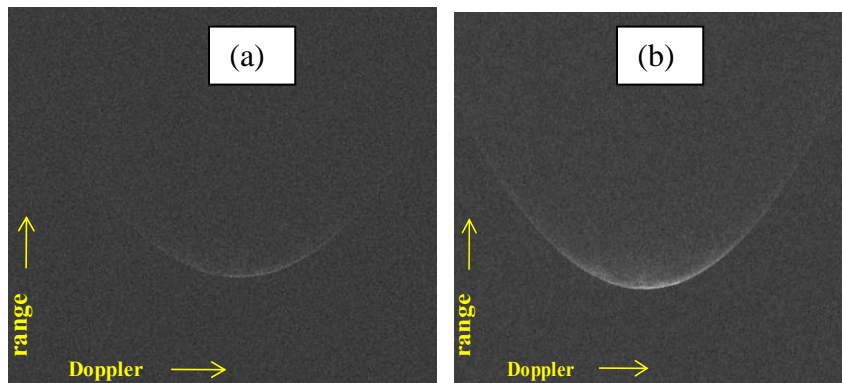


Fig 2. Doppler-delay images of Venus, taken on 2010 DOY-297, Goldstone Solar System Radar processing: a) single 34m antenna illumination; b) 2-34m antenna phased-array illumination, showing greatly improved image quality.

Figure 2 also demonstrates why an array is more reliable than a single dish system. For these observations, we were supposed to use 3-3m antennas at Goldstone. However, the day of the observing run, one of the 34s' m antennas had an anomaly. Even without that one dish, we saw an increase of 4 in the power which allowed us to still be successful albeit with reduced signal/noise (Vilnrotter et al. 2011). If we had used only a single dish radar system and had that antenna not been available, we would have been totally unsuccessful.

3.0 ADVANTAGES OF A MULTIPURPOSE FACILITY EMPLOYING UPLINK ARRAYING TECHNIQUES

- An **array is a more reliable** resource than a single dish. If the 70m is down for any reason, so too is the radar facility. The same is true for the high power klystron tubes used for the radar. At the time of this writing, one of the two tubes used to get to 450 kW at Goldstone is unusable. In addition, the 70m antenna that houses the NASA solar system radar was down for seven months under going depot level maintenance. During that time

NASA had NO ground radar capability whatsoever. However, with an array, if any given antenna is taken out for maintenance or is in an anomalous condition, little performance is lost. For example, losing a single antenna out of 25 would be a loss of only 2% of the array downlink capability and only 1% of the uplink capability. Hence, reliability of the array is more resistant and robust to operational “down time” or element failures.

- **Virtually 24/7 availability.** Whereas radar observations on the DSN 70m antenna comprise < 3% of the available antenna time, on a NEO-focused array, some 25-30 times more antenna time could be available and thus 25-30 times the number of sources can be observed in a given year. This will dramatically help NASA reach the goal of tracking and characterizing 90% of NEOs $\geq 140\text{m}$ by 2020.
- **Spectrum management is not an issue** with the array. Since the high power, coherently combined beam forms $\sim 200\text{ km}$ above the earth, the FAA EIRP limit will not be violated since the transmission from each individual antenna is below the limit thereby obviating the need for a time-consuming coordination among a large number of Agencies.
- The **range resolution** of a radar system is determined by the spectral bandwidth available. At X-band, the International Telecommunications Union has allocation 150 MHz. The Goldstone Solar System Radar uses on 40 MHz of that allocation leading to a range resolution of 375 cm. At Ka band, however, the primary allocation is 2.6 GHz, and with the secondary allocation, a total of 4 GHz bandwidth is available leading to a range resolution of 3.75 cm: two orders of magnitude improvement! In addition, we are exploring means of obtaining 1 cm range resolutions- without going to the highly weather dependent W-band (90 GHz).
- The **angular resolution** of the proposed array in a bistatic or multistatic mode with aspects in the western US and in Australia operating at Ka band (33-37 GHz) and used in a radio astrometric mode (measuring to 1/100 of the beam) has an angular resolution of 0.015 milliarcsec; the equivalent of 5 cm at GEO.
- **Scalability.** If still higher resolution or greater sensitivity is desired, additional antenna elements can be added. At roughly \$1.5M per antenna element, increased capability can be added quickly and at a low cost.
- **Extensibility** to Ka band. This would be unique to NASA and, in a single-dish equivalent, provide 16 times the angular resolution of the 70m radar system as well as significantly improved range and range-rate measurement.
- **Radio science experiments** are usually conducted by transmitting signals from the spacecraft past/through the target of interest to the ground. However, spacecraft transmitters, $\sim 20\text{W}$, limit the signal to noise ratio and hence the science results. Using a high power uplink from the ground to the target to the spacecraft and then downlinking the data via telemetry (ala New Horizons) can increase the S/N by ≥ 1000 . Science using traditional “downlink” measurement techniques will also be improved due to the higher sensitivity of the array.
- **Lower Life Cycle Costs:** The array also offers lowers life cycle costs. The comparative costs of the Deep Space Network with that of the VLA and VLBA are indicative. Consider the cost of the “wall power” required to run the uplink transmitters. For NASA’s Deep Space Network (DSN), each 34m antenna employs a 20 kW transmitter. If we put 750 W transmitters on each of 5 - 12m antennas, then $(5)^2 * 750 = 19\text{ kW}$; the same capability. The cost for the 20 kW transmitters and for five 750 W is essentially the same. However, they are also about 50% efficient. So, the 20 kW transmitters need 40 kW of wall power to generate this much power whereas the five 750 W transmitters also at 50% efficiency need only 7.5 kW of wall power to operate. The difference of 32.5 kW per transmission is significant. Each 34m DSN antenna transmits about 9000 hours/year. At $\$0.105/\text{kW hr}^1$, the cost savings is \$31K per antenna per year. If all 6 34m antennas were replaced with arrays and transmitters as above, the cost savings is \$180K/yr. Over a 20 year span at 3% inflation, this equates to a savings of \$3.7M. And there is a proposal to build still more 34m antennas for the DSN, so this number should be taken as a lower bound.

¹ (http://www.eia.gov/electricity/monthly/epm_table_grapher.cfm?t=epmt_5_6_a data as of 22 April 2014)¹

4.0 METHODOLOGY

4a. Mitigation of the Principal Error Contributors

1. Accommodation of circuitry, transmission line, and antenna shape variation
2. Differential beam steer phase due to dish to target line-of-sight variation (“predicts” in DSN parlance)
3. Mitigation of propagation phase variation due to tropospheric effects

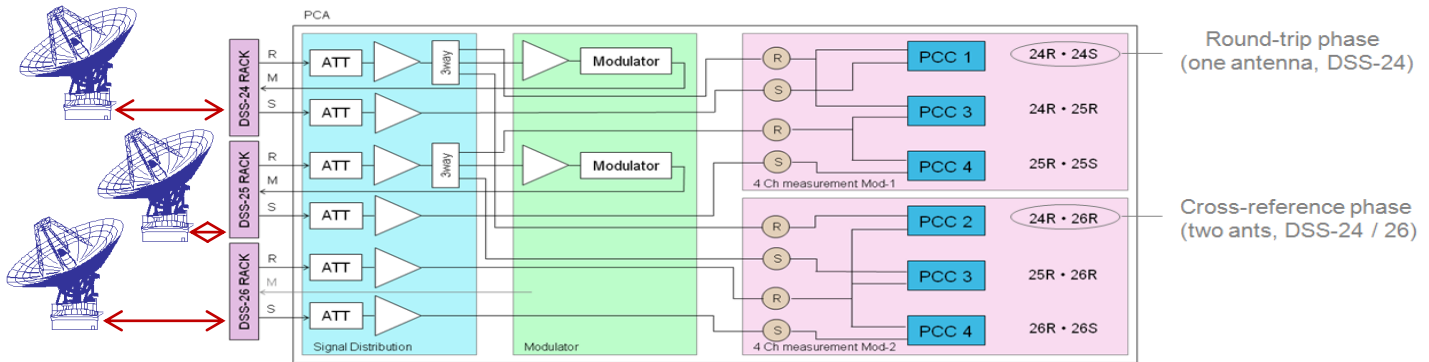
4b. Solution to Error Contributors

1. Continual Open-Loop Carrier Phase Control using Tx signal itself
Complex Envelope realization mitigates need for high precision time delay
2. Continual Adaptive Combining of antennas
Initial estimate of Antenna Reference Points (ARP) can be refined VLBI techniques
3. Transmit using phase estimates from a received signal
Expected to compensate for phase due to the atmosphere

In the next Figure, we show the method employed on 34m antennas to detect, measure, and compensate for circuit phase errors employed by NASA at the Jet Propulsion Lab.

PHASE COMPARATOR COMPENSATION & CONTROL AT SPC-10

- Phase Comparator, Phase Modulator, and Signal Distribution Assemblies
- Signal itself used as reference for round-trip and cross-phase measurements



- Cross-reference measurement determines exciter “wake-up” phase **AT ANY TIME**
- Round-trip measurement introduces 180 deg phase ambiguity, which must be resolved

Fig. 3. Continuous Open-Loop Phase Control at NASA Goldstone Signal Processing Center (SPC) 10 using the Tx Signal Itself: Mitigates Phase Errors due to: Transmission Lines, Circuitry, and Antenna Deformation

5.0 ATMOSPHERIC FLUCTUATION PHASE ERROR CORRECTION: RF ADAPTIVE OPTICS

The tropospheric contribution to phase errors can be significant and are certainly unpredictable as evidenced by water vapor radiometer measurements. However, there is no apparent linkage between meteorological data and the observed phase fluctuations. Thus, ground-level measurements are not accurate indicators of what is occurring higher in the atmosphere. Developing an all-encompassing, universal model to predict the phase stability of a particular site would be extremely difficult. Furthermore, if the longest baseline of the array is even a few kilometers in extent, different weather conditions above individual array elements would make precise modeling of

atmospheric phase fluctuations of the array as a whole an extraordinary challenge. Evidence for this is shown in Fig 4 for the Australia Telescope Compact Array in Narrabri, NSW. Whereas the gains of the inner five closely spaced (50-300 m) antennas relative to an array reference point are similar and track each other as a function of elevation, the gain for the more distant antenna 3 km distant falls off much more rapidly and dramatically.

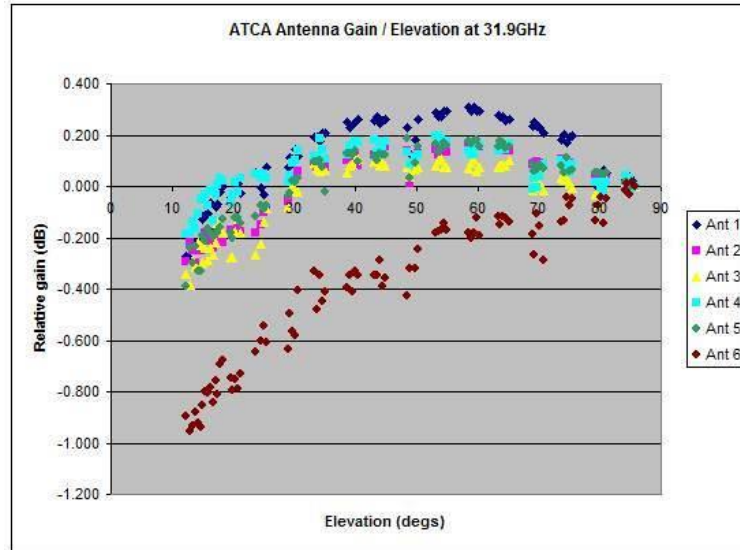


Figure 4: Measured gain as a function of elevation for the Australia Telescope Compact Array antennas at 31.9 GHz. Note that antenna 6 is 3 km distant from antennas 1-5.

A simpler solution is to develop a real-time atmospheric compensation methodology at radio frequencies just as the optical astronomers have done by using a sodium laser to get an artificial “star” in the telescope field of view simultaneously with the object of interest. The method developed for the open loop uplink arraying demo, the RF analog of the well known adaptive optics methodology, will be used to mitigate in real-time the varying atmospheric fluctuations. The scheme handles the general case where each dish-to-target path may be looking through different columns of atmosphere due to the wide spacing of the elements.

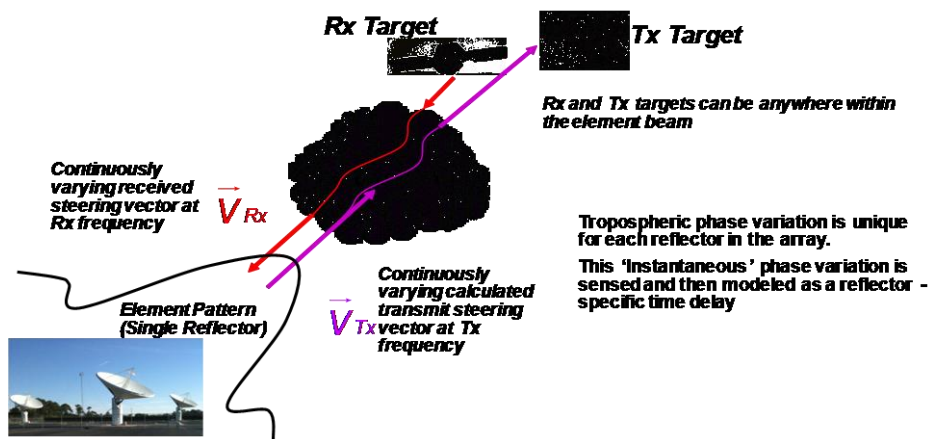
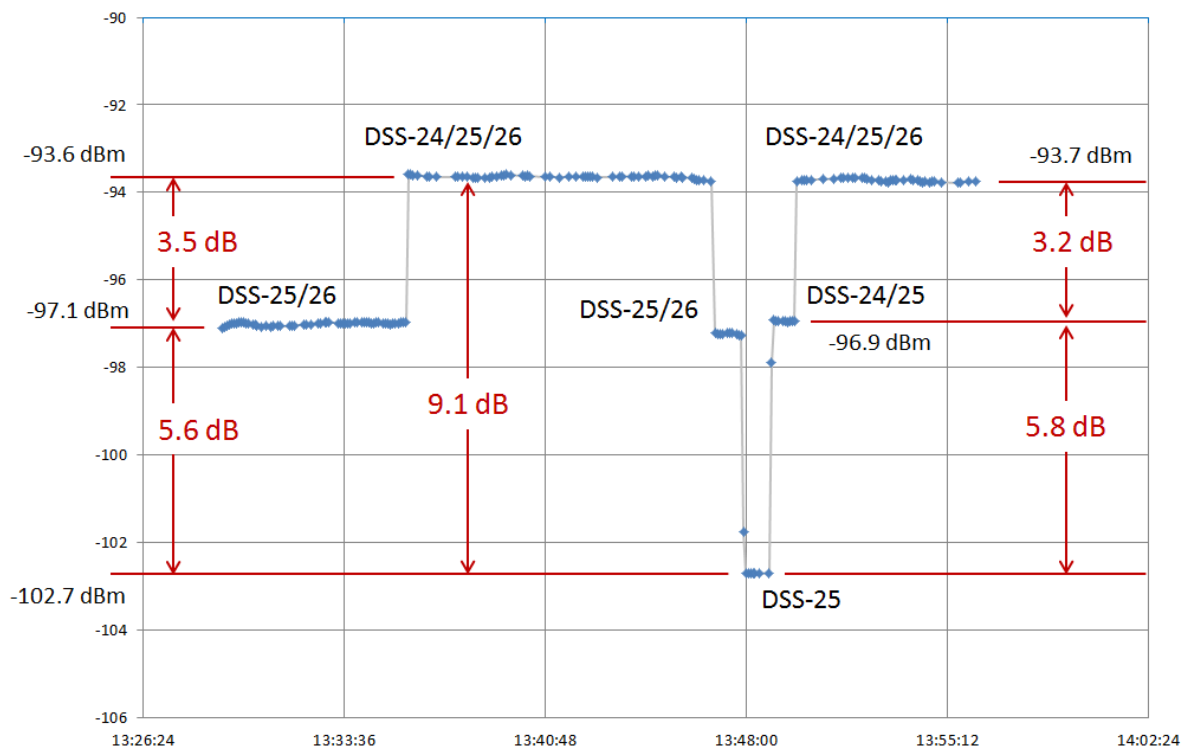


Fig. 6. Pictorial Representation of Tropospheric Fluctuation Phase Error Correction via RF Adaptive Optics

Real time propagation variations can be mitigated by using a signal from a source of known angular position- preferably within the primary beam- received through the varying atmosphere; assumed to be flat on average. Since the open loop receive circuit is operational, variations in the received steering vector from the flat atmosphere will measure the magnitude of the tropospheric phase contribution in the direction of the received signal. If the antenna beam width is small, then tropospheric variation toward the Tx target will be essentially the same as to the receive target. Since the effect due to atmospheric variation, mainly due to water vapor and turbulence, is one of frequency independent time delay, the total differential delay obtained from measured receive variation can be applied to calculations of the steering vector for the transmit direction. Any known astrometric source may be used for this purpose, such as another satellite, quasar or supernova remnant. For Ka band, the planets are excellent calibration and reference sources.

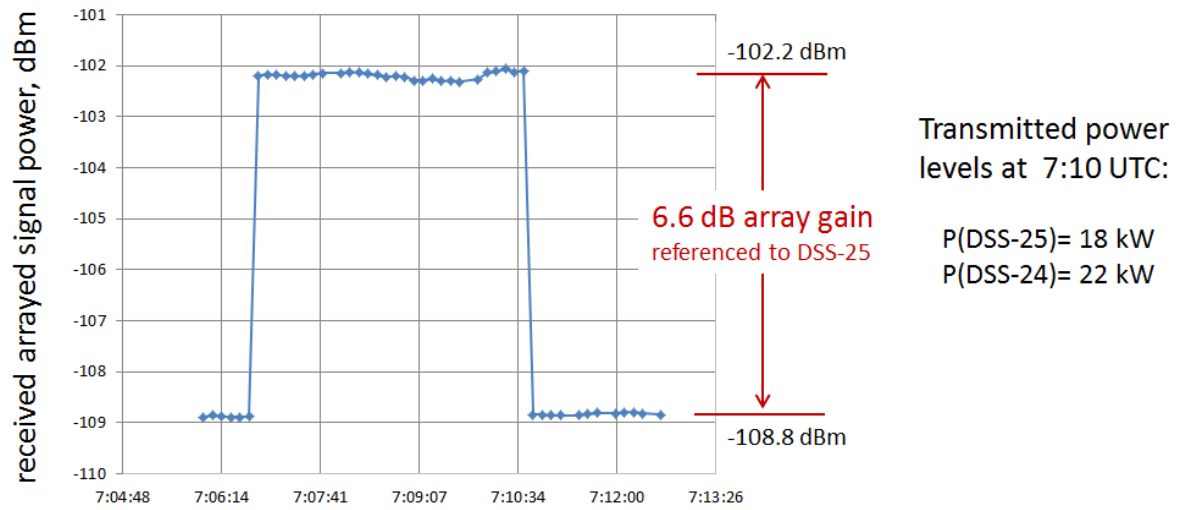
6.0 MEASURED RESULTS OF REAL TIME OPEN LOOP DEMONSTRATIONS

Since the proof resides in measured data, first we show the results of the expected EIRP gain from a coherently phased uplink array. In this monograph, we show the successful demonstration of the N^2 effect both for a spacecraft in cruise trajectory (EPOXI; see also V. Vilmrotter, D. Lee, T. Cornish, P. Tsao, L. PaaI and V. Jamnejad 2010) as well as for a spacecraft orbiting another planet (Mars Odyssey). The latter demonstration represents the first successful coherent uplink arraying to another planet and offers a technique for getting increased power to a spacecraft in case of emergency operations.



EPOXI Uplink Array track on DOY-179, 2008, showing two- and three-antenna array gains approaching 6 and 9.5 dB, as predicted by theory

Fig. 6. Measured EIRP increase for three arrayed 34m antennas at Goldstone through target satellite EPOXI



Portion of the first Odyssey (M010) Uplink Array track in 2014, showing 6.6 dB two-antenna array gain with a Mars-orbiting spacecraft, referenced to DSS-25 (lower transmitted power)

Figure 7: First demonstrated capability to coherently uplink signals from widely spaced antennas to a spacecraft orbiting another planet.

7.0 NEXT STEP: KaBOOM: Ka BAND OBJECTS OBSERVATION AND MONITORING



Figure 8: Overhead shot of the KaBOOM site at the Kennedy Space Center. It is comprised of 3-12m diameter antennas. The operations center is seen just to the right of center. Spacing between the antennas is 60m.



Fig 9. Current Array Configuration of Three 12m Reflector Antennas at the Kennedy Space Center

Going beyond X-band for NASA involves Ka band tracking capabilities. Although the demonstration of coherent uplink arraying at Ku band has been successful and gives us encouragement that such techniques will be successful at Ka band, we have no data to bear out that hypothesis. Furthermore, we have not yet shown how to make uplink arraying reliable in an operational sense because we have not previously had the resources to undertake that aspect.

In Phase 1, building on a previous deployment of a three-antenna element interferometer capability maturation facility to demonstrate the N^2 effect, where N is the number of antennas in the array, the antennas were moved to NASA's Kennedy Space Center to ensure physical and IT security as well as lower life cycle costs. The three-antenna element interferometer system will re-validate previously obtained X-band data, and reestablish the overall system baseline performance incorporating lessons learned from an initial implementation. Thereafter, we intend to demonstrate coherent uplink arraying with real-time atmospheric fluctuation correction at Ka band 30-31 GHz. All of these demonstrations involve the interferometer system in a communications mode.

In Phase 2, we shall convert the space communication system at KSC to a Ka band (34-36 GHz) radar system to demonstrate a high power, high resolution radar system. Our initial goal is 10 cm range resolution. If we can obtain Ka band transmitters with a bandwidth of 33-37 GHz, we may be able to attain a 10 cm range resolution. In addition, we are beginning to envision a technique to drive the range resolution to 1 cm! These are theoretical resolutions, and systematic errors etc. may increase these resolutions slightly.

In Phase 3, given successful performance and passing of NASA and our partners' funding decision gates, we shall construct a larger, multi-element array to increase capability. The Table below shows the potential advantages. It is evident that even with a modest number of antennas (15), there is the possibility of a 100% increase in maximum imaging distance over the current 70m capability and the ability to track objects over a volume more than 8.6 times larger than is current possible. Even a modest system of antennas can provide a substantially greater uplink power than is current available at Ka band. However, this assumes coherent uplink arraying at Ka-band can be successfully accomplished. Phase 1 is designed to demonstrate this capability and its limitations; i.e.- where does Ka band uplink arraying break down or become ineffective. This latter point is important. To date, none of the coherent uplink arraying demonstrations using widely separated antennas has failed! This is because time and funding constraints has not allowed any of the teams to test to failure. We plan to rectify this deficiency with KaBOOM. The beauty of an array system is that antennas and transmitters can be deployed as requirements, desirements and funding become available. The totally funding need not be allocated at once.

The advantages of operations at Ka band are enormous:

- For radar applications, an increased spectral bandwidth allocation of 4 GHz (vs. 40 MHz at X-band- Goldstone radar) thereby leading to a dramatic increase in spatial and range resolutions (5 cm vs. 400 cm) as well as more power on target (Table I). This is enabled by the newly developed 30 kW Ka band klystrons having a bandwidth of 4 GHz [33-37 GHz].
- For NEO, Space Situational Awareness, and orbital debris cases, the increased power afforded by uplink arraying radars can help to better characterize objects and can track and characterize objects farther out than current radars (Table I).
- For geolocation applications, measurement of atmospheric fluctuations coupled with other sensor data has the potential to increase the accuracy and precision of ground-based target location.
- For space communication purposes, the wider spectrum allocation (10X wider than at X-band) will allow for more data to be sent at a given time and compliments NASA’s on-going optical communications efforts.
- For radio science, the 100-1000x increase in possible uplink power will allow for more precise determination of planetary properties.

Table I. Comparison of Current and Proposed Systems

	# Antennas	Power (TW)	Maximum Distance (AU)
Current State of the Art	70m; X-band; 460 kW	1	11
	12m; Ka-band; 100 kW; 50% efficiency	1	0.9
		15	200
		25	525
		50	2150
		100	8600

The target satellite chosen for the Phase 1 Ka band demonstration, WGS 3, has an elevation as seen from KSC of only 10 degrees- meaning an air mass 5.6X greater than that toward the zenith. This constraint of increased attenuation and scintillation coupled with the non-ideal Ka band weather provides a highly challenging environment. Since this is a demonstration for NASA as well as for other partners who may deploy larger systems using these techniques in non-Ka band-pristine locations, we have deliberately chosen a difficult challenge.

7.0 APPLICABILITY TO NASA HUMAN SPACEFLIGHT AND SCIENCE ENDEAVORS

Part of the “flexible path” NASA is embarking on calls for exploration- robotic and crewed- to asteroids. Before a crew is sent to an asteroid, it is most likely, based on past exploration NASA practices that a robotic precursor will be sent to investigate. However, which asteroid will be chosen? If the “wrong” asteroid is chosen, we shall have lost or wasted a decade of time and funding. Hence it is incumbent upon us to select the best target.

It is well known that radar is an ideal technique to characterize (size, shape, spin, porosity) near earth objects and precisely determine their orbits (up to 5 orders of magnitude more precise than optical determinations). Radar measurements can prevent potential mission targets from being “lost.” Many NEO’s are lost shortly after discovery using optical techniques. However, radar observations can anchor the orbit of an object for decades or in some cases centuries. Furthermore, higher powers and thus farther distances can be achieved with an arrayed system thereby (a) expanding the search volume for NEO’s (a factor of ~150 for an array of 100 antennas), and (b) through characterization, narrow the potential target list thereby reducing the risk of sending a robotic precursor mission to the “wrong” asteroid.

Large arrays with high power transmitters on each antenna could lead to an NEO Early Warning System. In the next figure, we show the current capability, and what is possible with a large array: extending the area of tracking from 0.1 AU (1 Astronomical Unit is the average Earth-Sun distance, ~ 150M km).

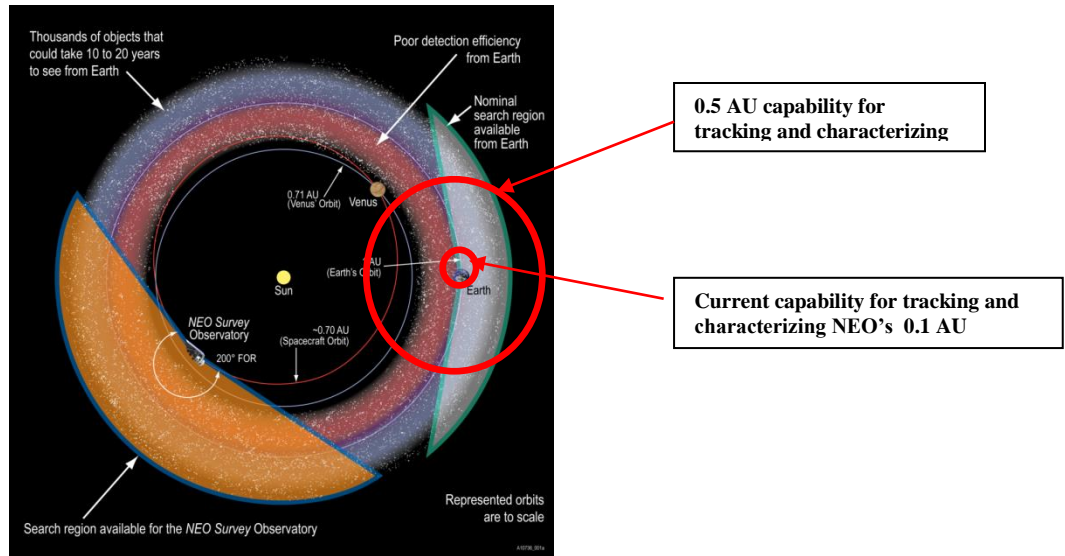


Fig 10: Possibility for a NEO Early Warning System

8.0 APPLICABILITY TO ORBITAL DEBRIS AND SPACE SITUATIONAL AWARENESS (SSA)

The National Space Policy states the following regarding Orbital Debris remediation (ODR): “For the purpose of minimizing debris and preserving the space environment for the responsible, peaceful, and safe use of all users, the United States Shall:

- Pursue research and development of technologies and techniques, through the Administrator of the National Aeronautics and Space Administration (NASA) and the Secretary of Defense, to mitigate and remove on-orbit debris, reduce hazards, and increase understanding of the current and future debris environment.”

NASA’s recognition of the importance of ODR technology coupled with limited resources provides the basis for guidance on ODR investments and activities. In summary:

- NASA related funded and unfunded work should extend only to technology development, not operational systems. NASA has no plans to establish an operational role in ODR.
- NASA will focus efforts in ODR technology development on Technology Readiness Level (TRL) 1-4 concepts.
- Current investments and activities should demonstrate nonduplicative cross-cutting relevance to the technology roadmap areas prioritized in the Agency’s Strategic Space Technology Investment Plan.

As time goes on, orbital debris has become and will continue to become and ever increasing source of risk to rocket launches, to the International Space Station, and to government and commercial space assets. Tracking of orbital debris on cm (or even mm) size scales and larger has become concomitantly more imperative. (Limiting Future Collision Risk to Spacecraft: An Assessment of NASA’s meteoroid and Orbital Debris Programs,” National Research Council report: The National Academies Press at http://www.nap.edu/catalog.php?record_id=13244). Here again, Goldstone has made a contribution- the statistics of the numbers of small particles. The Goldstone 70m

antenna can detect individual small particles, but the beam size is far too small to track these particles. The proposed array, with broad primary beam antennas, has the advantage. This type of system can complement and supplement the activities of the Space Fence. However, although NASA's Goldstone antenna can detect individual small (< 1 cm) particles, they are not tracked because the particles move through the very narrow antenna beam (2.2 arcmin) too quickly. Furthermore, that antenna is so busy tracking NASA spacecraft, that little time, perhaps only 100 hours/year, are available for such observations.

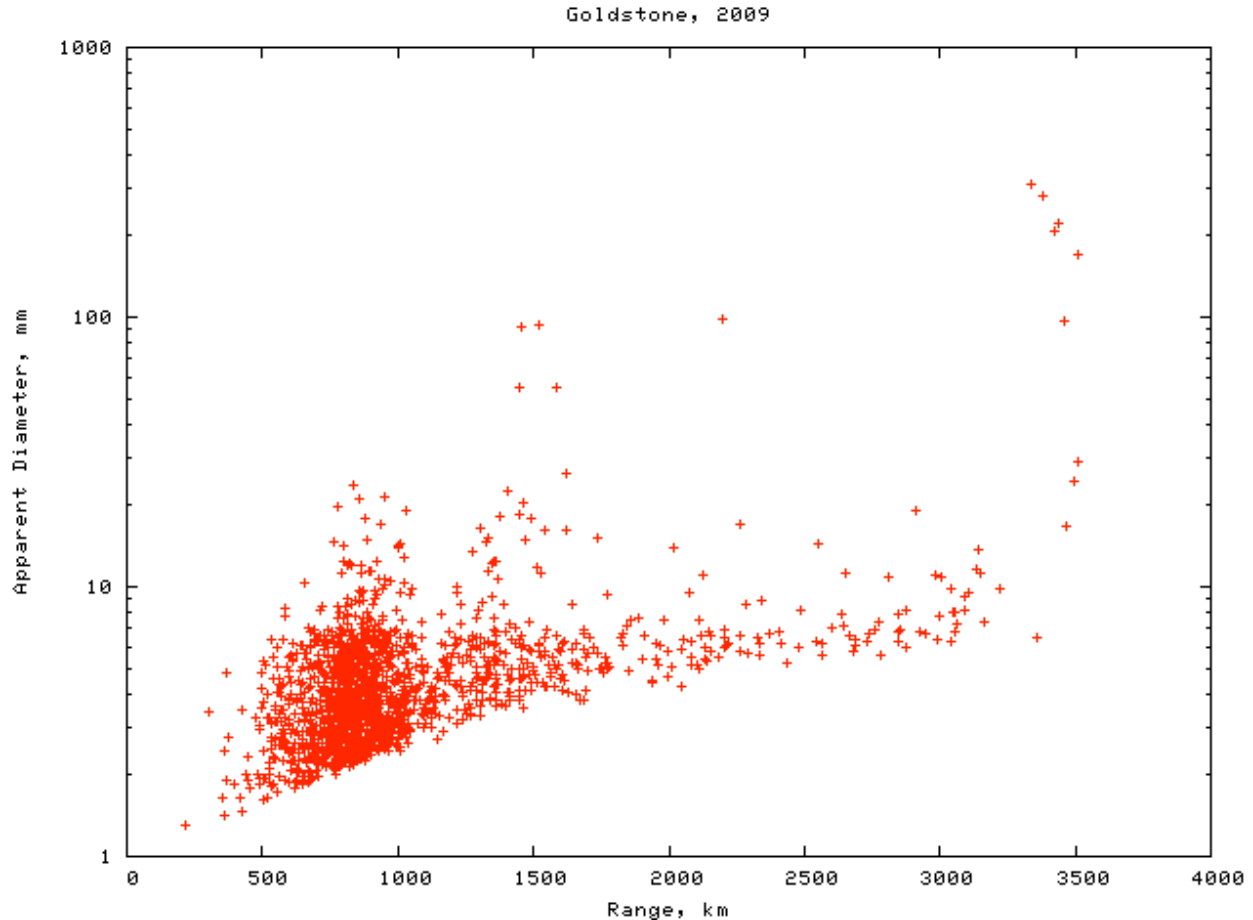


Fig. 11. Goldstone Orbital Debris Radar Apparent Diameters detected in FY2009. Each cross represents an individual particle.

For SSA, the recent GAO report : “SPACE ACQUISITIONS : Development and Oversight Challenges in Delivering Improved Space Situational Awareness Capabilities” (GAO-11-545 May 2011: <http://www.gao.gov/new.items/d11545.pdf>) summed up the current status: “The United States’ growing dependence on space systems for its security and well-being—such as for missile warning; intelligence, surveillance, and reconnaissance; communications; scientific research; weather and climate monitoring; and positioning, navigation, and timing—makes these systems vulnerable to a range of intentional and unintentional threats. These threats range from adversary attacks such as anti-satellite weapons, signal jamming, and cyber attacks, to environmental threats such as harsh temperatures, radiation, and collisions with debris and other man-made or natural objects, which have been increasing rapidly over the past several years.”

Detailed simulations of the capabilities of a phased array radar system indicates that current gaps in the following areas of SSA knowledge can be addressed

- **Detect/Track/Identify:** Uncued detection, Unexpected maneuvers
Discriminate between closely spaced objects

- **Characterization:** Orbital Debris, Satellite break-ups, collisions
- **Threat Warning and Assessment:** Conjunction assessment, Re-entry prediction

A Ka band system using coherent uplink arraying techniques and bistatic and multistatic radars can meet and probably exceed the goals or, at the very least, compliment a 90 GHz system. Specifically, at Ka-band range resolution of 5 cm and a spatial resolution (using a US-Australia baseline) of ~ 5 cm can be achieved.

The first demonstration of using coherent uplink arraying in a radar mode has been undertaken at the Jet Propulsion Lab (Vilnrotter and Tsao) using two 34m beam wave guide antennas at X-band. Precise calibration of Array Radar antenna phase required, and the phase calibration can be accomplished via the “Moon-Bounce” method. In addition, closed-loop phase control is required to maintain phase calibration since temperature variations and equipment instabilities degrade coherence. The proposed array radar approach provides simultaneous projection of OD velocity vectors onto three independent baselines, thus enabling trajectory determination from a single Array Radar observation!

In brief, three 34m BWG antennas with 20 kW transmitters at 7.18 GHz are available. The array spans ~ 500 meters with antenna null-to-null beamwidths ~ 170 mdeg and spacings: DSS-24 – DSS-25 baseline ~ 23 mdeg and DSS-24 – DSS-26 baseline ~ 15 mdeg. For a single antenna: EIRP of 34m antenna, 20 kW transmitter; two antenna array: peak EIRP of 34m antenna, 80 kW transmitter; three antenna array: peak EIRP of 34m antenna, 180 kW transmitter. CAVEAT: Simultaneous multi-frequency/multi-baseline operation remains to be shown. This would enable processing of simultaneous echoes from different baselines.

The receiver operates at 1 mega-samples/sec with 8 bits quantization. For the array radar signal processing, we developed algorithms to process received samples, and display the data. Event frequency and amplitude, spectrograph displays were generated. Several sources were observed, including Atlas 5 and Cosmos booster stages, and event frequency, amplitude, and spectrograph displays were generated as shown in Fig. 12.

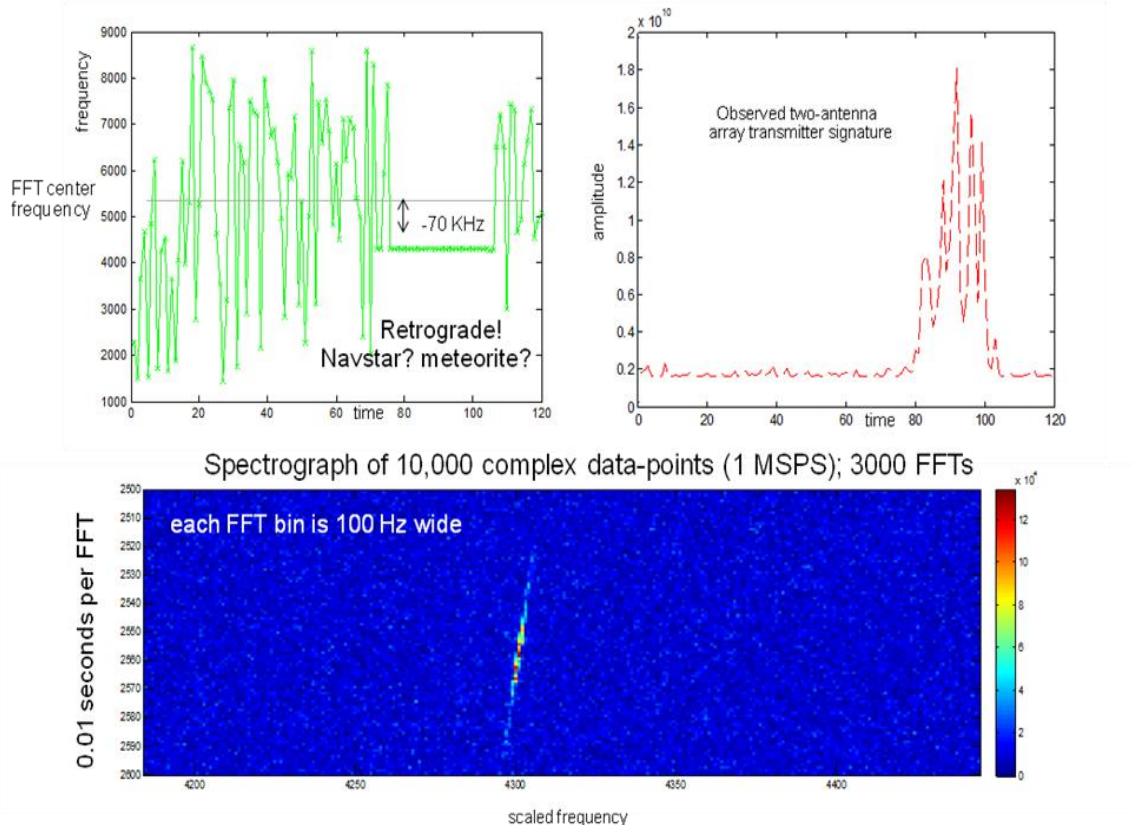


Fig 12. Preliminary Experimental Results of a two antenna radar array

10.0 APPLICABILITY TO MAPPING OUT “NO DRIVE” ZONES ON MARS

A radar stealth region $2000 \times \sim 300$ km was discovered on Mars (Muhleman et al. 1993). It was dubbed a stealth region because there were virtually no radar reflections from the area. The best postulate for the stealth property is that the soil has a density of 0.5 gm/cc thus inhibiting reflections. It is believed that the Mars rover Spirit got stuck in such a “sandy” constituency- away from the stealth region (McCustion, private communication 2012). If we could map the entire planet and the locations of such regions- e.g., “no drive” zones, we might save or increase the useful lifetimes of multibillion dollar assets.

Once we have established that Ka band uplink arraying is possible, the next step is to establish a large array with high power Ka band transmitters. Using arrays with radio *astronomy* resolution techniques on a US-Australia baseline at Ka band would provide a ground resolution of 50 km. Useless for attaining our goal. Using radio *astrometry* techniques where we can measure to 1/100 of a beam or better, we might expect to get resolutions of 0.5 km. Interesting, but the fundamental limitation is the diameter of the Earth. With ~ 100 kW transmitters on a feasible number of antennas, we can use the radar array in a SAR mode and let baseline be determined by the motion of the planets. Preliminary calculations show that at Mars closest approach to Earth, about $5.46 * 10^7$ km, a *ground resolution of 10 m can be obtained in one hour of observation!*

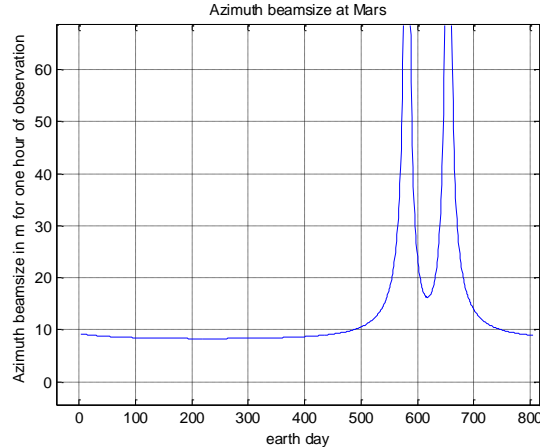


Fig 13: Estimate synthesized beam azimuth resolution

The problem of $1/r^4$ losses from radar observations still exists, though. If, however, NASA puts an orbiter around Mars with Ka band receive capability, a high power signal from Earth could reflect off the planet and be captured on the spacecraft thereby changing a $1/r^4$ diminution to merely a $1/r^2$ effect while not requiring a high power Ka band transmitter on the orbiter.

11.0 APPLICATION TO BEAM SAILING

Here, a microwave beam will strike a solar sail delivering electromagnetic pressure. This is a natural extension of solar sailing in that the microwave beam assists the sun (Benford and Benford 2003). A qualitative study of how much mass could be delivered to Mars at closest approach was undertaken in 2004, and provided intriguing results (Benford and Benford 2004). NASA is exploring a proof of concept demonstration as a prelude to a large scale, high power uplink array as one of the tools in the human space exploration tool kit.

We have tentatively identified several cubesat missions that will employ solar sails as targets of opportunity to test the beam sailing ideas. In particular, the demonstration will be carried out in a manner similar to on-off photometric astronomical measurements. The cubesat with sail deployed will cruise under solar radiation pressure. At an appropriate time, the coherent microwave beam from several antennas will be aimed at the cubesat, the acceleration measured, and then beam will be turned off. Repeating this sequence should provide the evidence we seek. Note that we are not at this time looking for an ablative inner portion of the sail as Benford and Benford (2003) describe.

Additionally, there are two dependencies to undertake a successful demonstration: (1) we must first demonstrate “uplink very long baseline interferometry” (see figure 15) since initially we shall employ existing assets at NASA’s Goldstone tracking stations. The N2 effect of the 70m antenna plus 3 or 4 3m antennas should provide the power required, and (2) we shall need an inertial measurement unit with sufficient sensitivity and resolving power to measure the anticipated small push from the combined microwave beams.

12.0 GOING FORWARD PLANS

As of this writing, only KaBOOM is funded. However, we are making detailed cost estimates for replacing the communication mode system in KaBOOM with a purely Ka band radar system: KARNAC [Ka band Array Radar for NEO Accurate Characterization]. The transformation from KaBOOM to KARNAC – Plan A has been jump started by our obtaining three surplus Ka band TWTs (35-37 GHz spectral response, each 50 kW peak power).

KARNAC has the potential to transmit the same power as the Goldstone radar system (450 kW) albeit with a range resolution of 10 cm as opposed to the 400 cm at Goldstone.

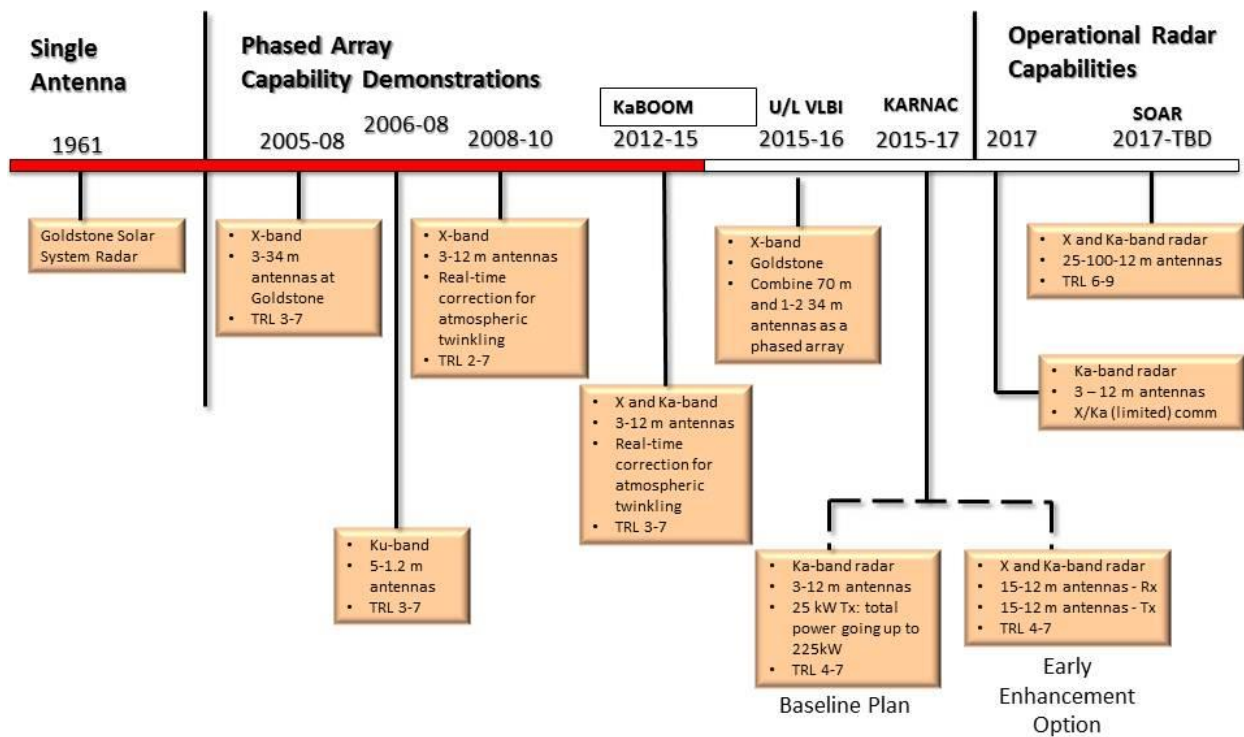


Fig. 14 Maturation and Capability Demonstration Roadmap

13.0 CONCLUSIONS

We have the funding in place for a demonstration of uplink arraying at Ka band with real-time compensation for atmospheric fluctuations. We shall start at X-band in a space communications mode (Fall 2014), next repeating the demonstration but at Ka band (30-31 GHz), and finally demonstrate Ka band radar capabilities (late CY 2015). Additionally, a skunk works project team has been established: (1) minimal formal management- all team members will manage the project together; (2) a top-drawer technical team. Finally, it is gratifying that all the hard work and accomplishments over the last decade on various aspects of uplink arraying has garnered broad support with NASA and with our mission partners. When a myriad of folks said “You’ll NEVER be able to send commands to a space craft via a coherently phased uplink array,” the teams just went ahead and successfully overcame all the challenges. The teams have never been phased by having to follow the phases through all the phases of the demonstration, thereby accomplishing what NASA does so well: Doing the “impossible” every day

14.0 ACKNOWLEDGEMENTS

It is a pleasure to thank Pete Aragona, Jason Crusan, Lindley Johnson, Richard McGinnis, Bill Marinelli, Michael Miller, Chris Moore, and Marc Seibert, for their guidance and encouragement regarding the demos, MIT/LL both at the lab in Lexington MA and at Kwajalein, Marshall Islands for their guidance and patience with the lead author- a simple radio astronomer and Les Johnson at MSFC for beam sailing discussions. The author also is appreciative of NASA management at Headquarters at the Division, Directorate, and Agency levels for allowing these demonstrations to be undertaken and for the use of the facilities.

REFERENCES

Benford, G. and Benford, J. 2003, The Planetary Report, January/February

Benford, G. and Benford, J. 2004, private communication.

Geldzahler, B., 2011, “Coherent Uplink Arraying Techniques for Next Generation Space Communications and Planetary Radar Systems,” SPIE Defense, Security + Sensing

Geldzahler, B., 2011, “Coherent Uplink Arraying: A Self-Calibrating, Nearly Stand-Alone System,” Military Sensing Symposium/Broad Area Maritime Surveillance

McCouston, private communication 2012.

Muhleman et al. 1993, Science , 253, 1508.

Vilnrotter, V., Lee, D., Cornish, T., Tsao, P., PaaI, L. and Jamnejad, V. IEEE A&E SYSTEMS MAGAZINE, MAY 2010

V. Vilnrotter, P. Tsao, D. Lee, T. Cornish, J. Jao, M. Slade 2011, Proceedings of the IEEE Aerospace Conference, Big Sky, Montana, March 2011